

The EUSO-TA detector: status and performance

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EUSO-TA is an on-ground telescope, located at the Telescope Array (TA) site in Black Rock Mesa, Utah, USA. The main aim of the instrument is observation of Ultra High Energy Cosmic Rays through detection of ultraviolet light generated by cosmic-ray showers. EUSO-TA consists of two, 1 m² Fresnel lenses with a field of view of about 11° × 11°. Light is focused on the Photo Detector Module composed of 36 Hamamatsu multi-anode photomultipliers, for a total of 2304 channels. The telescope is housed in a shed located in front of the TA Fluorescent Detector. We present the results from 5 observational campaigns of EUSO-TA, including detected cosmic rays and meteors. Also shown are the studies done with dedicated artificial light sources, such as LEDs and lasers.

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1. Introduction

JEM-EUSO is a space-borne mission dedicated to detecting cosmic rays of the highest energies [1]. It will observe the ultra-violet light from Extended Air Showers (EAS) generated by cosmic rays in the atmosphere with a Fresnel lens based optics and super-fast single photon counting camera. Unlike the existing, on-ground experiments, JEM-EUSO due to its distance from the surface of the Earth, would observe a much larger volume of the atmosphere, significantly increasing the number of collected events and, hopefully, localising their sources in the Universe.



Figure 1: EUSO-TA (front right), EUSO-SPB [2] (front middle) and TAFD (back) (photography by M. Mustafa)

EUSO-TA is a fully functional ground telescope (fig. 1) in EUSO family. It is located at Black Rock Mesa, Utah, at the site of one of the fluorescence light detectors of the Telescope Array (TA) experiment [3]. From there it observes, simultaneously with TA, artificial light and cosmic ray events, allowing for tests of the technology, calibration of the detector and reduction of the systematic uncertainties of the measurements. The location and pointing allows for observation of TA's Central Laser Facility (CLF) and Electron Light Source (ELS) (fig. 2).

2. EUSO-TA instrument

The 1 m^2 squared EUSO-TA Fresnel lenses are fabricated from UV transmitting polymethylmethacrylate (PMMA). The lenses focus light on the $17 \text{ cm} \times 17 \text{ cm}$ PDM, composed of 36 Multi-Anode Photomultiplier Tubes (MAPMTs) [4] each containing 64 anodes, for a total of 2304 pixels (fig. 3). Four MAPMTs form an EC-Unit, each with a dedicated Cockcroft-Walton based High Voltage Power Supply. The working voltage of the phototubes is 1000 V, and is automatically reduced in case of too bright light on an EC-Unit in single microseconds.

Each PMT is read out by one SPACIROC1 ASIC, which are distributed onto 6 EC-ASIC boards [5]. The time resolution of ASICs is $\sim 30 \text{ ns}$. This implies a saturation at about 28 counts

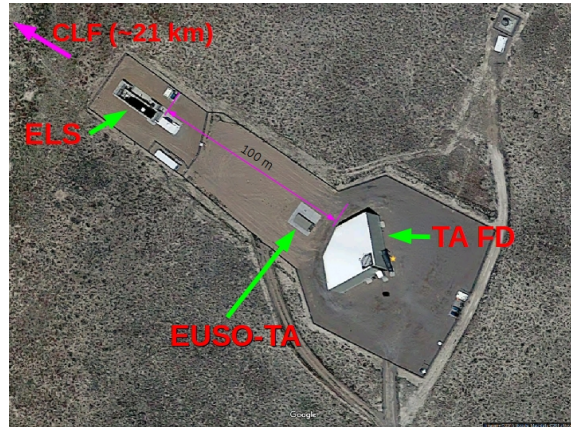


Figure 2: The position of EUSO-TA on the Telescope Array Fluorescent Detector Site in Black Rock Mesa, Utah, USA. The EUSO-TA dome is placed just in front of the TAFD. In its line of sight there is Electron Light Source and Central Laser Facility instruments of TA. The photograph was taken from Google Maps.

on the frame – the Gate Time Unit (GTU) – of $2.3 \mu\text{s}$. The frame is followed by 200 ns of dead time.

The digitized counts from 6 EC-ASIC boards are read into a ring buffer of the PDM board. The buffer contains 128 GTUs and is read out on request, which can be made with an external or internal trigger. The data is then transferred to a Cluster Control Board (CCB), which contains a 2nd-level trigger board [6] and CPU board. Together with Clock board (CLKB), GPS board, house keeping board and low voltage power supply [7] they form a Data Processing (DP) unit. The exchange of information is made with encapsulated packets. The amount of information contained in a packet increases with the level of processing, the final packet stored by the CPU on a hard disk containing counts from ASICs, additional information from the PDM board, CCB, CLKB, GPS, and so on [8].

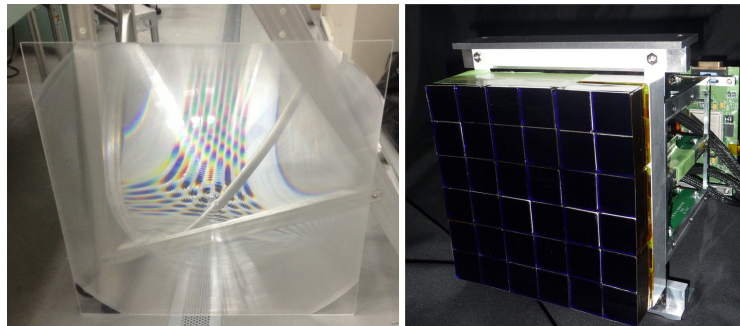


Figure 3: Front lens of EUSO-TA (left) and the PDM array and the front-end ASIC boards mounted on the PDM frame (right).

3. EUSO-TA campaigns

The EUSO-TA lenses and mechanical structure was installed in Telescope Array Fluorescence Detector site in Black Rock Mesa, Utah, USA in March 2013. EUSO-TA PDM and readout system has been installed in February/March 2015. Since then, there were 5 observational campaigns of EUSO-TA, 4 in 2015 and 1 in 2016. More than 136 hours of observation was performed using TAFD external trigger, thus could contain UHCER events.

4. Results

4.1 Stars observations

EUSO-TA can observe stars up to magnitude in Johnson’s blue filter $M_B \simeq 6.5$ on sums of 1280 frames (about 3.2 ms observation time). This limiting magnitude is very approximate, since it depends largely on atmospheric conditions and stars spectra (EUSO-TA spectral window is far from standard Johnson filters). An example of 1280 stacked frames with a few stars clearly visible is shown in fig. 4. The Hipparcos catalogue [9] is superimposed on the image for the 4 brightest stars, allowing us to recognize Algol and stars with $M_B \geq 5.5$.

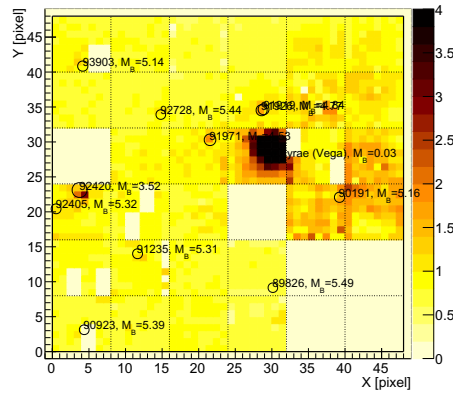


Figure 4: Sum of 1280 frames acquired with EUSO-TA with Hipparcos catalogue brightest stars’ positions superimposed. The colour scale denotes the brightness of each pixel in arbitrary units after flat fielding.

The stars can be used as point sources to analyze the point spread function (PSF) of our detector. Initial analysis performed with a fit of standard gaussian profile gives a PSF with average FWHM of 2.35 pixels, close to the edge of the frame, where it is the biggest. It is well within the requirements for UHECR showers observations. Stars were also used to determine the field of view of the instrument, which is $10.6^\circ \pm 0.3^\circ$.

4.2 Other “slow” signals

While EUSO-TA is designed for observations of millisecond-scale events, it can observe phenomena happening on much longer timescales. The most numerous are flashes from airplanes and sunlight reflected by satellites. EUSO-TA has also observed meteors, as can be seen in fig. 5. Such

observations are unique due to their time resolution. It could be important in a case of detecting a phenomenon similar in appearance – strangelets [10].

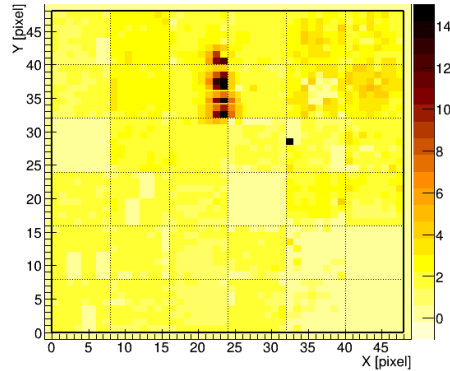


Figure 5: A meteor track detected by EUSO-TA. The picture shows an overlap of four averages of 1280 frames. The color scale denotes the uncalibrated detector counts.

4.3 Lasers observations

To study the EUSO-TA response to a known light source we have used the light coming from the TA’s CLF, distant from EUSO-TA by about 21 km. The CLF shoots vertically a laser of 355 nm wavelength in front of the detectors [11].

During standard observation nights the CLF is shot every half an hour for 30 s with 10 Hz shooting frequency. The scattered light of the ~ 3 mJ beam is clearly visible traversing through the EUSO-TA field of view on 6 to 8 frames, depending on the shot and acquisition time synchronization. The spot length is 6-8 pixels depending on the position on the frame (fig. 6, left), which is consistent with expectations. The registered light intensity is dependent both on the atmospheric conditions – obstruction of the light path due to clouds can be easily noticed – and the fine features of the detector’s focal surface.

In addition to the CLF we have also performed measurements of a Global Light System (GLS) laser - a mobile UV laser of Colorado School of Mines. The laser can be shot with energies in the range of about 1–86 mJ, with pointing adjustable in two dimensions. The mechanics featured automatic changing of the pointing, allowing for easy “swipes” through the field of view and, additionally, emulating the orbital condition where the CR increases its distance to the detector with time.

Figure 7 shows reconstructed brightness of the laser track in the detector vs laser shot energy, for the distance of 33 km. In the tested energy range of 4–22 mJ, which is outside the saturation region, the dependency is linear, showing that the detector behaves as expected. For the lowest energies of 2– ~ 3 mJ, only the few brightest laser events were reconstructed, resulting in a brightness cut on the data.

4.4 UHECR

We consider so far 4 as detected and 4 as candidate UHECR events (fig. 8). Their distances

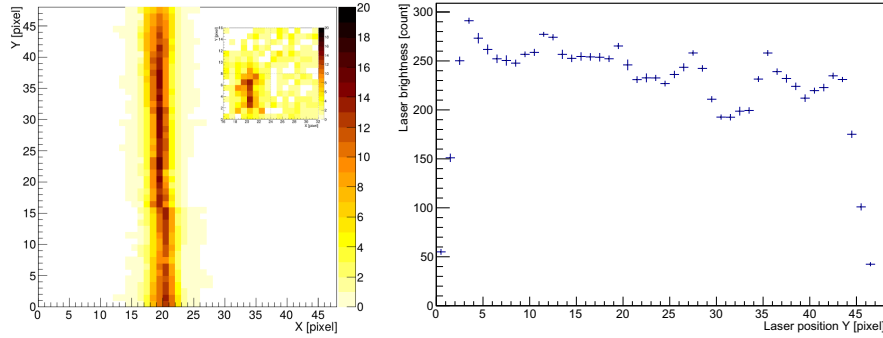


Figure 6: Left: an average of 259 tracks of CLF laser. The subfigure shows a zoomed part of a single frame containing the laser. The colour scale denotes the uncalibrated counts. Right: a profile histogram of 259 CLF tracks. Each point value shows an average of summed counts of all tracks at a corresponding vertical position of shots. The brightness is lower for spots falling into space between PMTs or out of the PDM.

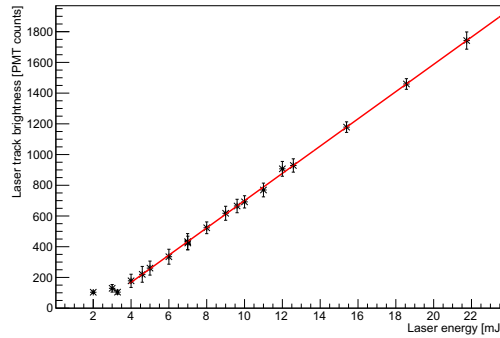


Figure 7: Reconstructed brightness of the GLS laser track vs its energy for the distance of 33 km from the detector. The plot is joined from 2 shooting sessions, altogether encompassing energy range of 2–22 mJ. Each point was calculated from a few dozen shots. Fitted line shows good linearity of the detector response in the range 4–22 mJ, while for lower energies only the brightest tracks were reconstructed, enforcing a low brightness cut on the data points.

from the detector vary between approximately 1.5 and 9 km, while the energy between $10^{17.7}$ – $10^{18.8}$ eV, according to TA measurements. An example event is shown in fig. 9.

These first events registered with EUSO technology allowed us to vastly improve our reconstruction and simulation algorithms, the work still being in progress. However, EAS parameters had to be derived from TAFD which, thanks to larger FoV and higher time-resolution, could see the shower movement. The closer look at the parameters of UHECR that happened in EUSO-TA field of view during its operation starts to reveal the detection capabilities of the detector.

5. EUSO-TA future plans

EUSO-TA is the best available way to test technology for existing and future EUSO family

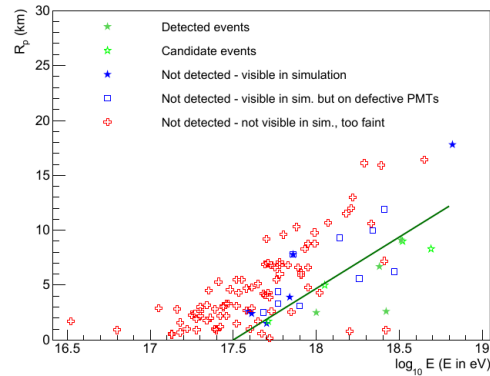


Figure 8: All UHECR registered by TAFD in the EUSO-TA field of view during its operation. Markers' shape and colour denotes status in EUSO-TA: not detected, detected and detected in simulation, with a distance-energy sensitivity border emerging from the data.

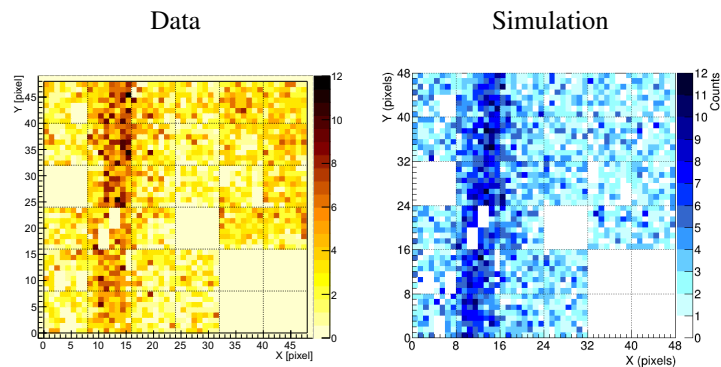


Figure 9: Example of an UHECR observed by EUSO-TA. Left shows the real data in photoelectron counts, right the simulation. The shower had an energy of $10^{18.36}$ eV and was 2.6 km from the telescope (data from TA measurements).

experiments, as it allows for stable field observations for extended time periods. However, the observation time could be significantly increased if we switched from manual to remote automatic operation of the instrument. This would eliminate a need of a shifter on-site and associated costs, in principle allowing to collect data in all the observational periods during the year. This task will require introduction of some additional mechanics and control devices, as well as a non-remote automatic working test period.

Additionally, we want to develop an autonomous trigger for slow events such as meteors, strangelets or lightnings, which are mostly discarded by the TAFD. This would expand the scientific goals of EUSO-TA.

6. Summary

EUSO-TA employs a new technology of observing cosmic rays, with use of Fresnel lenses and multi-anode photomultipliers. The detector has registered a 4 confirmed and 4 candidate UHECR

during its five observational campaigns, proving that this approach works well for the main goal. Additionally, a number of “slow” events such as stars, meteors and planes have been observed allowing for an extension of scientific objectives.

The main goal of the detector, however, was the test of hardware capabilities and stability. It proved invaluable in the modifications applied for EUSO-SPB and Mini-EUSO [12] detectors. In near future we plan an upgrade and automatization of the telescope to increase the sensitivity and duty cycle.

Acknowledgments

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